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A HIGH-RESOLUTION STUDY OF THE HYDRA A CLUSTER WITH *CHANDRA*: COMPARISON OF THE CORE MASS DISTRIBUTION WITH THEORETICAL PREDICTIONS AND EVIDENCE FOR FEEDBACK IN THE COOLING FLOW

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ABSTRACT

The cooling flow cluster Hydra A was observed during the orbital activation and calibration phase of the *Chandra Observatory*. While the X-ray image of the cluster exhibits complex structure in the central region as reported in McNamara et al., the large-scale X-ray morphology of the cluster is fairly smooth. A spectroscopic analysis of the ACIS data shows that the gas temperature in Hydra A increases outward, reaches a maximum temperature of 4 keV at 200 kpc, and then decreases slightly at larger radii. The distribution of heavy elements is nonuniform, with a factor of 2 increase in the Fe and Si abundances within the central 100 kpc. Beyond the central 100 kpc the Si-to-Fe abundance ratio is twice solar, while the Si-to-Fe ratio of the central excess is consistent with the solar value. One of the more surprising results is the lack of spectroscopic evidence for multiphase gas within the bulk of the cooling flow. Beyond the central 30 kpc, the ACIS spectra are adequately fitted with a single-temperature model. The addition of a cooling flow component does not significantly improve the fit. Only within the central 30 kpc (where the cooling time is less than 1 Gyr) is there spectroscopic evidence for multiphase gas. However, the spectroscopic mass deposition rate is more than a factor of 10 less than the morphologically derived mass accretion rate at 30 kpc. We propose that the cooling flow region is convectively unstable owing to heating by the central radio source, which significantly reduces the net accretion rate. In addition, we show that the mass distribution within the central 30–200 kpc region scales as $\rho_d \propto r^{-1.3}$, intermediate between an NFW and a Moore profile, but with a best-fit NFW concentration parameter ($c_{\text{NFW}} = 12$) approximately 3 times greater than that found in numerical simulations. However, given the limited photon statistics, we cannot rule out the presence of a flat-density core with a core radius less than 30 kpc.

Subject headings: cooling flows — galaxies: clusters: individual (Hydra A) — intergalactic medium — X-rays: galaxies

1. INTRODUCTION

The launch of the *Chandra X-ray Observatory* (Weisskopf et al. 2000) provides astronomers with the first opportunity to view clusters of galaxies with the same spatial resolution as ground-based optical telescopes. This improved resolution is especially beneficial in the study of cooling flow clusters in which most of the interesting astrophysics is confined to the central 100 kpc (Fabian & Nulsen 1977; Sarazin 1986; Fabian 1994). One of the great mysteries concerning cooling flows has always been the contrasting implications of the X-ray observations with observations at other wavelengths. While X-ray observations suggest that a substantial amount of hot gas (up to $1000 M_{\odot} \text{ yr}^{-1}$) should be condensing out of the ambient medium within an extended region of radius 100 kpc, activity at all other wavelengths is constrained to the central 10 kpc (e.g., Hu, Cowie, & Wang

1985; Heckman et al. 1989; McNamara & O’Connell 1992; Donahue et al. 2000). With *Chandra*’s subarcsecond resolution, even scales of 10 kpc can be resolved in clusters at all redshifts.

Some of the first results from *Chandra* show that the central regions of cooling flow clusters are not simple spherically symmetric systems, but are morphologically and spectroscopically very complex. *Chandra* has shown that radio emission from the central galaxy in cooling flows can have a significant impact on the X-ray morphology of the central regions. X-ray emission has been detected from the radio hot spots and the central point source in 3C 295 (Harris et al. 2000; Wilson, Young, & Shopbell 2000). Spectroscopic analysis of the *Chandra* data shows that the emission from the hot spot is probably due to synchrotron self-Compton and is not produced through an interaction between the relativistic plasma in the radio lobes and the thermal gas. The X-ray morphology of the central region of the Hydra A cluster (Abell 780) reveals the presence of cavities in the X-ray-emitting gas coincident with the radio lobes (McNamara et al. 2000). The *Chandra* A observation of Hydra A also shows the presence of a heavily absorbed point source at the core of the radio galaxy (McNamara et al. 2000; Sambruna et al. 2000). The column density of absorbing material derived from spectral analysis of the central point source agrees with that observed in H I absorption (Taylor 1996). The *Chandra* observation of the core of the Perseus cluster shows a very complex X-ray

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morphology produced by interactions between the hot gas and the radio-emitting plasma and also low energy absorption produced by an infalling gas rich galaxy (Fabian et al. 2000).

A massive cooling flow was reported in Hydra A based on *Einstein* (David et al. 1990), *ROSAT* (Peres et al. 1998), and *ASCA* (Ikebe et al. 1997) observations with estimated mass deposition rates ranging from 60 to 600 $M_{\odot} \text{ yr}^{-1}$. *Einstein* MPC and *ASCA* spectral analyses give an emission weighted temperature of ~ 4 keV (David et al. 1990; Peres et al. 1998). Temperature profiles have been derived from the *ASCA* data by Ikebe et al. (1997), Markevitch et al. (1998), and White (2000). The *ASCA* observation of Hydra A also reveals the presence of an Fe abundance gradient (Ikebe et al. 1997; Finoguenov, David, & Ponman 2000) with a radially increasing silicon-to-iron ratio indicating a predominance of Type II supernovae enriched material at large radii. From the *Chandra* observations we are able to spectroscopically determine the mass deposition rate, gas temperature, and abundance of heavy elements at a much greater spatial resolution than previously possible.

In the first paper on Hydra A, we concentrated on the interaction between the central FR type I radio source (3C 218) and the ambient cluster gas. In this paper we concentrate on the spectroscopic and morphological properties of the cooling flow region, derive high-resolution temperature and abundance profiles of the gas within the central 250 kpc, and determine the total gravitating mass distribution in the core of the cluster. In addition, we compare the observed mass distribution in Hydra A with theoretical expectations from numerical simulations of dark matter halos (e.g., Navarro, Frenk, & White 1997; Moore et al. 1999).

The remainder of this paper is organized as follows. In § 2 we discuss the large-scale X-ray morphology of the Hydra A cluster, the surface brightness profile, and the gas density distribution. In § 3 we present the spectroscopic results concerning the temperature and abundance profiles of the hot gas. The total gravitating mass and gas mass distributions are derived in § 4 along with a comparison with theoretical predictions. The roles of weak and strong shock heating by the central radio source in producing a convectively unstable region within the cooling flow are discussed in § 5, and the main results are summarized in § 6. We assume $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3$, and $\Omega_{\Lambda} = 0.7$ throughout the paper. At a redshift of 0.0538, the luminosity distance to Hydra A is 240 Mpc, and $1'' = 1.05 \text{ kpc}$.

2. X-RAY MORPHOLOGY AND SURFACE BRIGHTNESS

Hydra A was observed during the orbital activation and calibration phase of the *Chandra* mission with both ACIS-I (1999 October 30) and ACIS-S (1999 November 2). Because of flares in the background-charged particle flux, the data were screened to exclude all time intervals when the background rate was more than 20% higher than the quiescent rate. The resulting screened integration times are 24,100 and 18,950 s for the ACIS-I and ACIS-S observations. Only the standard grade set (0, 2, 3, 4, 6) was used in our analysis. The gain was corrected using the CIAO task *acis_process_events* with the most recent version of the gain table appropriate for observations taken at a focal plane temperature of -110° C (*acisD1999-09-16gainN0004.fits*). This task also corrects for charge transfer inefficiency (CTI) effects in the front illuminated (FI) chips.

In McNamara et al. (2000), we showed a high-resolution image of the combined ACIS-I and ACIS-S observations of Hydra A to illustrate the interaction between the radio lobes and the ambient cluster gas. To illustrate the large-scale X-ray morphology of Hydra A we use only the ACIS-I observation because of the larger field of view and lower background. A background ACIS-I image was generated by co-adding observations taken at a focal plane temperature of -110° C in the *Chandra* public archive after removing point sources from the observations. The same gain corrections and background screening criteria were applied to the observations, resulting in a total integration time of 73 ks for the background image. The background image was then mapped onto the sky with the same aspect solution as in the Hydra A observation. The background-subtracted image was corrected for detector and vignetting effects using an exposure map generated by convolving the aspect histogram with the mirror effective area and ACIS-I QE maps. Finally, the background-subtracted and exposure-corrected image was smoothed with a $\sigma = 4''$ Gaussian. Contours of the smoothed ACIS-I image are shown in Figure 1 overlaid on the optical image of the cluster. The complex X-ray morphology within the central region discussed in McNamara et al. (2000) is evident in Figure 1. Beyond the central region, the X-ray morphology of Hydra A is fairly smooth with slightly elliptical contours, indicating that the cluster has not experienced a major merger in the recent past. The X-ray isophotes are aligned in the same southeast-northwest direction and have the same ellipticity as the *I*-band isophotes shown in McNamara (1995). There is thus a remarkably strong alignment between the stellar and gaseous components in the Hydra A cluster from scales of a few kiloparsecs up to 300 kpc.

While the ACIS-I observation is useful for illustrating the large-scale two-dimensional morphology of Hydra A, we use the ACIS-S observation to derive an X-ray surface

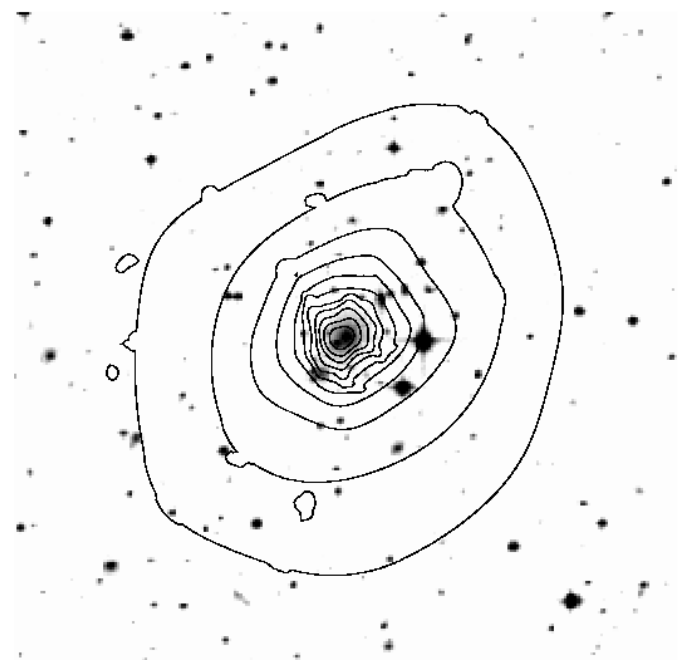


FIG. 1.—ACIS-I contours of the background-subtracted and vignetting- and exposure-corrected ACIS-I data overlaid on the optical image of the Hydra A cluster. The figure is 8:53 (540 kpc) on a side.

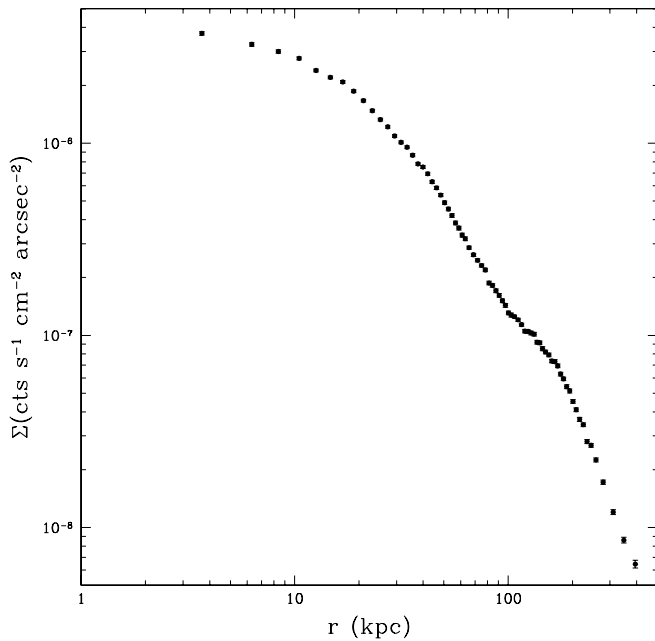


FIG. 2.—0.3–7.0 keV surface brightness profile of the Hydra A cluster from the ACIS-S observation. The data are background-subtracted, vignetting- and exposure-corrected, and binned to produce 2000 net counts per annulus.

brightness profile owing to the better photon statistics on the back-illuminated (BI) S3 chip. The BI chips on ACIS have higher QEs than the FI chips for photon energies below 4 keV. A broadband (0.3–7.0 keV) surface brightness profile derived from the S3 data is shown in Figure 2. A similar procedure was used to perform the background subtraction as discussed above for the ACIS-I data by co-adding 106 ks of ACIS-S3 data taken at a focal plane temperature of -110° C. The same reductions were applied to this background image, as discussed above for the

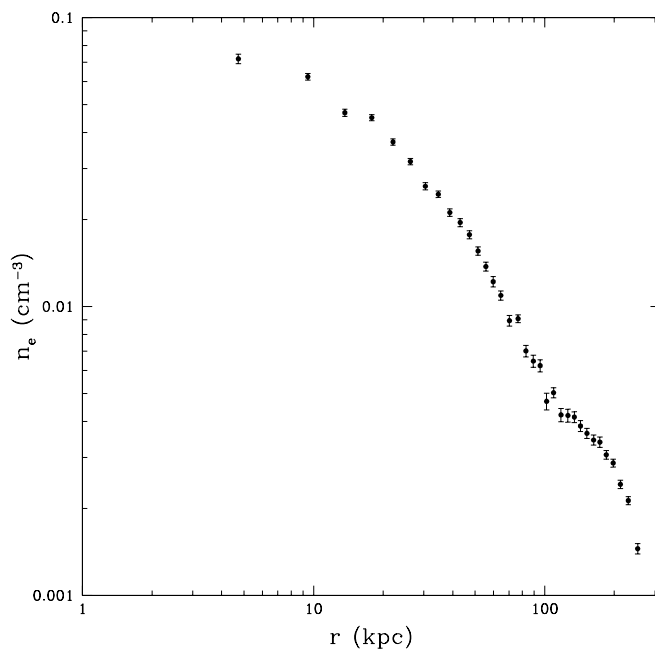


FIG. 3.—Deprojected electron-density profile of the Hydra A cluster with errors derived from 100 Monte Carlo simulations.

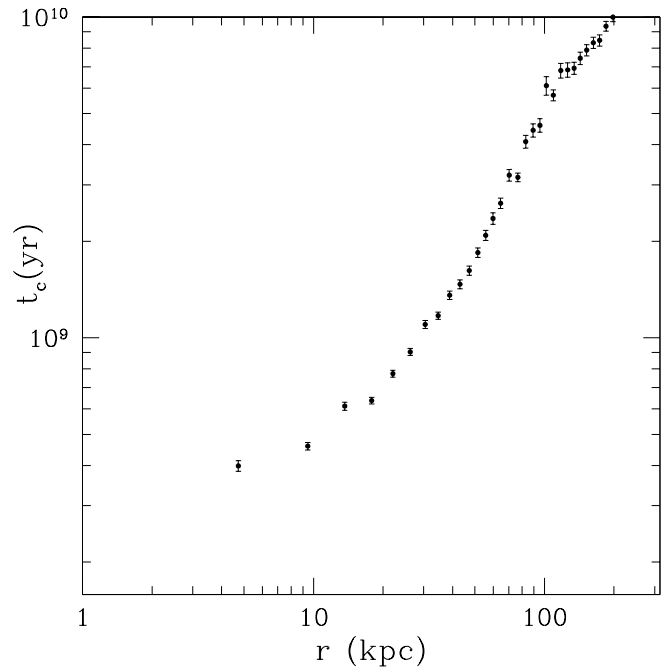


FIG. 4.—Isobaric cooling time of the hot gas as a function of radius

ACIS-I background image. Background count rates were then obtained from the same annuli as the total count rates. The origin in the surface brightness profile corresponds to the location of the central point source. The annuli were chosen to ensure approximately 2,000 net counts per annulus. *Chandra*'s high resolution shows that there are two inflection points in the surface brightness profile at approximately 100 and 200 kpc. It is obvious from Figure 2 that neither a single- nor a double- β model will provide a good fit to the data. The break in the surface brightness profile at 100 kpc is also evident in the *ROSAT* PSPC observation discussed by Mohr, Mathiesen, & Evrard (1999). Beyond the central 200 kpc, the surface brightness is approximately characterized by a power law with a slope corresponding to $\beta = 0.7$, in good agreement with the *Einstein* IPC results (David et al. 1990) and the *ROSAT* PSPC results (Mohr et al. 1999; Vikhlinin, Forman, & Jones 1999).

Owing to the complexity of the surface brightness in the central region of Hydra A, we use a deprojection technique to determine the gas density profile rather than fitting a multivariate function. The deprojection technique first determines the density of the gas in the largest annulus using the observed count rate and an assumed emissivity and then determines the density at progressively smaller radii after removing the projected emission from gas at larger radii. The fraction of the cluster emission in the outermost annulus arising from gas at larger radii is corrected using the best-fit β model to the PSPC observation of Hydra A in Mohr et al. (1999). We perform the deprojection 100 times using a Monte Carlo technique that randomizes the surface brightness profile based on the observed errors. The emissivity is assumed to be that of a 3.5 keV gas with a 40% solar abundance of heavy elements. In § 4 below we show that the gas temperature varies between 3.0 and 4.0 keV in this region and that the abundance varies between 0.2 and 0.6 solar. The assumption of a spatially independent gas emissivity only introduces a 3% uncertainty in the gas density. The mean gas densities and rms deviations from the

100 Monte Carlo simulations are shown in Figure 3. This figure shows that the density profile gradually steepens from $\rho_{\text{gas}} \propto r^{-0.4}$ at 30 kpc to $\rho_{\text{gas}} \propto r^{-1.6}$ at 100 kpc, and there is no obvious flat density core.

The isobaric cooling time, t_c , of the gas is shown in Figure 4. Note that t_c is less than a Hubble time ($\sim 10^{10}$ yr) within the central 200 kpc, and less than 1 Gyr within the central 30 kpc. This inner radius is important for the discussion below and appears to be the outer boundary of an isentropic core of multiphase gas.

3. SPECTRAL ANALYSIS

The ACIS-S observation of Hydra A produced 160,000 source counts within the central 250 kpc on the S3 chip. Spectra were extracted in eight concentric annuli centered on the point source with widths varying from $10''$ to $30''$ to ensure 20,000 source counts per spectrum. Spectra also were extracted with 40,000 net counts each, since better photon statistics are required to determine the Si distribution. We excised the emission from the central point source using a $2''$ radius circle. Background spectra were extracted from the same regions in the background fits event list and were normalized by exposure time. All spectra were extracted in PI space and rebinned to a minimum of 20 counts per bin. A single PI response matrix file (rmf) was generated using mkrmf and the FEF file appropriate for the aim point on S3, node 1, at a focal plane temperature of -110° C. Area files were generated for each spectrum based on the mean off-axis angle of the extracted photons in each spectrum. We fitted each spectrum to a variety of XSPEC models: an absorbed single-temperature model, an absorbed single-

temperature plus cooling flow model, an absorbed single temperature plus intrinsically absorbed cooling flow model, and an absorbed two-temperature model (see Tables 1–3). Only data between 0.5–8.0 keV were included in the analysis. Below 0.5 keV the detector gain is nonlinear, which is not yet properly treated in the response matrices. Above 8.0 keV, the charged particle background dominates the source emission.

The ACIS-I data were extracted in a similar fashion; however, several modifications are required to account for the nonuniformity of the FI chips due to radiation-induced CTI effects. Instead of a single rmf, a photon-weighted rmf was generated for each spectrum. In addition, the CTI damage to the ACIS-I chips produces a change in quantum efficiency at high energies (events far from the chip readout experienced significant grade migration, transforming acceptable grades to unacceptable grades) that we corrected using the ACIS-I QE maps at $T = -110^\circ$ C.

The spectral fitting results for the inner three annuli of the ACIS-S data are shown in Tables 1–3 and Figures 5–7. The spectrum from the inner $2''$ – $17''$ is the most complex, and significant reduction in χ^2 can be obtained by adding multiple spectral components. Table 1 shows that including a cooling flow component in the inner spectrum improves the fit with a reduction in χ^2 of 5 to 8 depending on whether N_{H} is fixed to the galactic value or allowed to vary. However, there is not an obvious improvement when simply examining the best-fit of the two models to the data (see Fig. 5). The spectral analysis in the innermost annulus where the cooling time is $\sim 10^8$ yr, shows that less than 24% (at the 90% confidence limit) of the total emission arises from gas

TABLE 1
COMPARISON OF SPECTRAL FITS FOR THE $2''$ – $17''$ ANNULUS

Model	kT (keV)	N_{H} (10^{20} cm $^{-2}$)	Z (solar)	\dot{M} (M_{\odot} yr $^{-1}$)	N_{HI} (10^{20} cm $^{-2}$)	kT_{low} (keV)	χ^2/dof
WABS*MEKAL	$2.78^{+0.08}_{-0.09}$	4.88	$0.36^{+0.06}_{-0.05}$	250/181
	$2.88^{+0.18}_{-0.11}$	$3.65^{+0.79}_{-0.80}$	$0.41^{+0.07}_{-0.06}$	242/180
WABS*(MEKAL	$3.00^{+0.15}_{-0.13}$	4.88	$0.44^{+0.06}_{-0.06}$	$8.9^{+4.1}_{-4.9}$	242/180
+ MKCFLOW)	$3.11^{+0.14}_{-0.15}$	$3.63^{+0.57}_{-0.68}$	$0.49^{+0.07}_{-0.07}$	$6.3^{+4.3}_{-4.3}$	237/179
WABS*(MEKAL	$3.05^{+0.11}_{-0.17}$	4.88	$0.44^{+0.06}_{-0.06}$	$8.7^{+4.3}_{-3.9}$	< 3.2	...	241/179
+ ZWABS*MKCFLOW)	$3.10^{+0.14}_{-0.15}$	< 2.3	$0.48^{+0.08}_{-0.06}$	29^{+17}_{-17}	$38.0^{+4.3}_{-3.9}$...	232/178
WABS*(MEKAL	$4.08^{+3.9}_{-1.02}$	4.88	$0.31^{+0.06}_{-0.05}$	$1.93^{+0.68}_{-0.32}$	240/179
+ MEKAL)	$3.72^{+2.4}_{-0.90}$	$3.56^{+0.80}_{-0.5}$	$0.37^{+0.06}_{-0.07}$	$1.91^{+0.59}_{-0.72}$	233/178

NOTE.—The results of fitting the ACIS-S data to a variety of spectral models (defined in column 1), kT is the gas temperature in single-temperature fits and the highest temperature in two-temperature fits, N_{H} is the hydrogen column density, Z is abundance of heavy elements relative to the solar value, \dot{M} is the mass deposition rate in models with a cooling flow component, and kT_{low} is the lower gas temperature in the two-temperature fit. When fitting the single temperature plus cooling flow model, the abundance of heavy elements is linked between the two components and the lowest gas temperature is fixed at 0.02 keV. Two rows are shown for each model. The top row shows the results with N_{H} fixed to the galactic value and the second row shows with results with N_{H} free. All error bars are shown at the 90% confidence level for one interesting parameter ($\chi^2_{\text{min}} + 2.71$).

TABLE 2
COMPARISON OF SPECTRAL FITS FOR THE $17''$ – $29''$ ANNULUS

Model	kT (keV)	N_{H} (10^{20} cm $^{-2}$)	Z (solar)	\dot{M} (M_{\odot} yr $^{-1}$)	N_{HI} (10^{20} cm $^{-2}$)	χ^2/dof
WABS*MEKAL	$2.97^{+0.11}_{-0.13}$	4.88	$0.35^{+0.06}_{-0.05}$	249/190
	$3.15^{+0.14}_{-0.15}$	$3.44^{+0.66}_{-0.67}$	$0.41^{+0.06}_{-0.06}$	236/189
WABS*(MEKAL	$3.15^{+0.17}_{-0.17}$	4.88	$0.38^{+0.06}_{-0.05}$	$7.6^{+4.3}_{-4.2}$...	244/189
+ MKCFLOW)	$3.24^{+0.18}_{-0.17}$	$3.65^{+0.72}_{-0.70}$	$0.42^{+0.07}_{-0.06}$	$4.8^{+4.4}_{-4.4}$...	236/188
WABS*(MEKAL	$3.15^{+0.17}_{-0.15}$	4.88	$0.38^{+0.06}_{-0.05}$	$7.4^{+4.5}_{-3.2}$	< 3.1	244/188
+ ZWABS*MKCFLOW)	$3.25^{+0.22}_{-0.18}$	$3.32^{+1.0}_{-2.3}$	$0.42^{+0.07}_{-0.06}$	$5.8^{+19.8}_{-5.2}$	< 4.2	236/187

NOTE.—See the notes for Table 1.

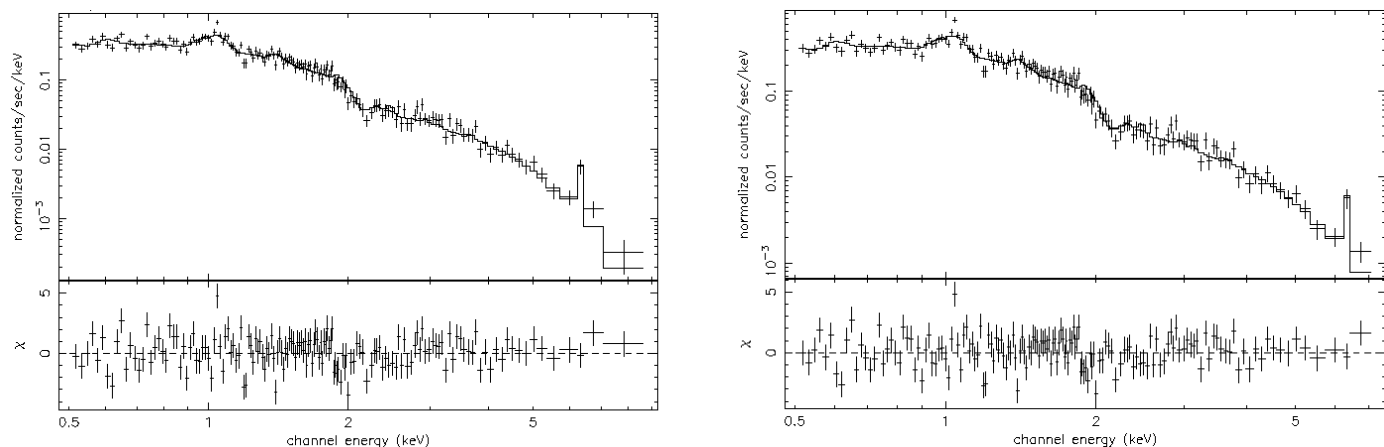


FIG. 5.—*Left*: Best-fit single-temperature model to the innermost ACIS-S spectrum ($2''$ – $17''$). *Right*: Best-fit single temperature plus cooling flow model to the innermost ACIS-S spectrum.

condensing out of the flow. While the absorbed cooling flow model produces the minimum χ^2 and largest \dot{M} , the best-fit galactic N_{H} in this model is more than a factor of 2 below the Stark et al. (1992) value. The intrinsic and galactic absorption are strongly correlated in the fitting process, and it is difficult to distinguish between them when fitting the spectrum of a low-redshift, low- \dot{M} cluster. Fixing N_{H} to the galactic value significantly reduces \dot{M} and limits the intrinsic absorption to less than $3.2 \times 10^{20} \text{ cm}^{-2}$ (see Table 1). Adding a cooling flow component to the second annulus only provides a slight improvement to the fit (see Table 2), while adding a cooling flow component to the third annulus

does not produce any improvement (see Table 3). Figures 6 and 7 show how well the spectra from the second and third annuli are fitted with a single-temperature model. Beyond the central 30 kpc, all of the spectra are well fitted with single-temperature models and the addition of more free parameters does not significantly improve the fits.

The lack of a strong spectroscopic signature for a massive cooling flow is not due to the small spatial binning of the spectra. In Figure 8 we show a spectrum extracted from the central 100 kpc, excluding the central point source, fitted to a single temperature plus cooling flow model with \dot{M} fixed at $140 M_{\odot} \text{ yr}^{-1}$ (the value obtained from a deprojection of the *ROSAT* PSPC data by Peres et al. 1998 after adjusting

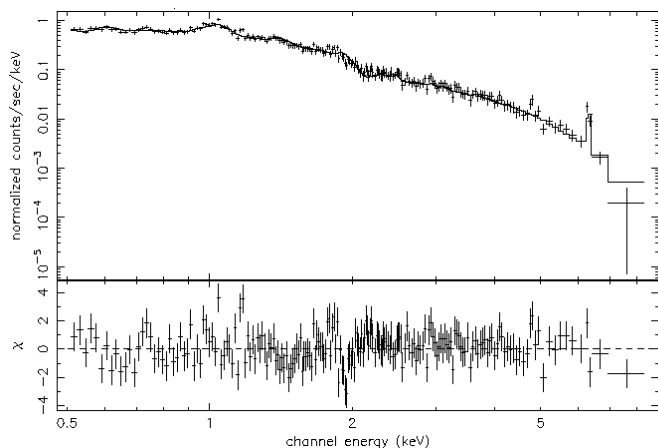


FIG. 6.—Best-fit single-temperature model to the ACIS-S spectrum extracted from an annulus between $17''$ and $29''$ from the central point source.

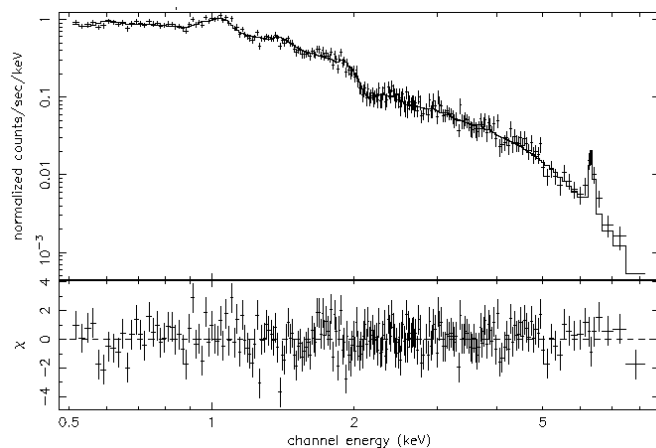


FIG. 7.—Best-fit single-temperature model to the ACIS-S spectrum extracted from an annulus between $29''$ and $43''$ from the central point source.

TABLE 3
COMPARISON OF SPECTRAL FITS FOR THE $29''$ – $43''$ ANNULUS

Model	kT (keV)	N_{H} (10^{20} cm^{-2})	Z (solar)	\dot{M} ($M_{\odot} \text{ yr}^{-1}$)	χ^2/dof
WABS*MEKAL	$3.25^{+0.12}_{-0.12}$	4.88	$0.34^{+0.06}_{-0.05}$...	259/232
	$3.36^{+0.14}_{-0.15}$	$4.03^{+0.70}_{-0.67}$	$0.38^{+0.06}_{-0.06}$	—	255/231
WABS*(MEKAL	$3.39^{+0.18}_{-0.17}$	4.88	$0.36^{+0.06}_{-0.05}$	<9.9	256/231
+ MKCFLOW)	$3.43^{+0.18}_{-0.17}$	$4.22^{+0.72}_{-0.71}$	$0.38^{+0.07}_{-0.06}$	<8.1	254/230

NOTE.—See the notes for Table 1.

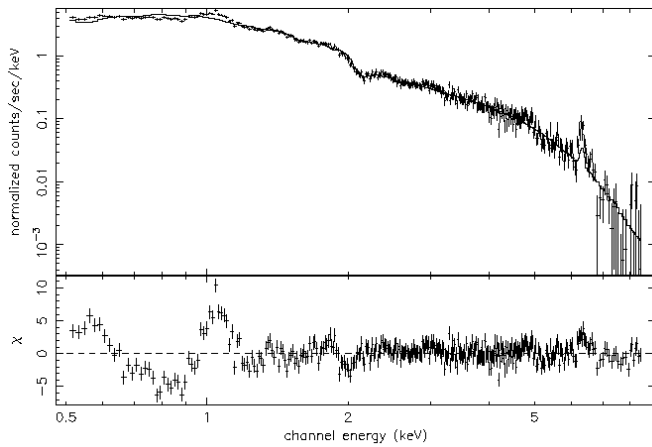


FIG. 8.—Best-fit of the ACIS-S spectrum of the central 100 kpc (excluding the central point source) to a single temperature plus cooling flow model with \dot{M} frozen at $140 M_{\odot} \text{ yr}^{-1}$ and all other parameters free.

to our adopted cosmology). All other parameters are treated as free parameters. Figure 8 shows that there are very large residuals below 2 keV and the resulting $\chi^2/\text{dof} = 1175/355$. Our results are in reasonable agreement with the spectral analysis of the *ASCA* SIS data on Hydra A by Ikebe et al. (1997) who obtained $\dot{M} = 30 \pm 15 M_{\odot} \text{ yr}^{-1}$ (scaled to our adopted cosmology). The implications of these results for cooling flows will be discussed below.

3.1. Temperature Profile

The gas temperature profile derived from the ACIS-S and ACIS-I data is shown in Figure 9. The inner two temperature points for the ACIS-S data were derived from the best-fit to a single absorbed temperature plus cooling flow model. The remaining temperature points for the ACIS-S data were obtained without a cooling flow component in

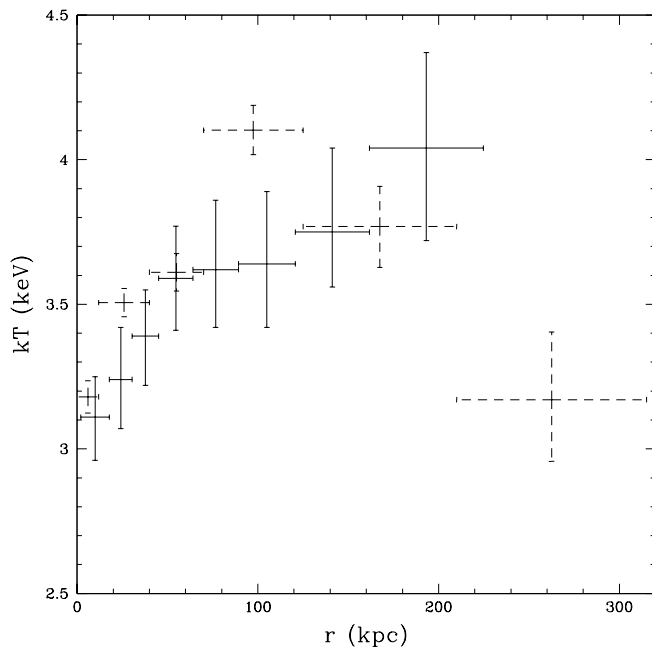


FIG. 9.—Gas temperature profile of the Hydra A cluster. The solid (dashed line) crosses show the spectral analysis results of the ACIS-S (ACIS-I) data. Error bars are shown at the 90% confidence level for one interesting parameter.

the spectral analyses. All ACIS-I spectra were fitted to single-temperature models with N_{H} fixed to the galactic value and a variable abundance of heavy elements. Figure 9 shows that there is good agreement between the two ACIS detectors and that the gas temperature in Hydra A increases from 3 to 4 keV within the central 200 kpc, as noted in McNamara et al. (2000). It is very difficult to obtain any temperature estimates from the ACIS-S data beyond 200 kpc owing to the high background in the S3 chip; however, the ACIS-I data show that the temperature drops to about 3 keV at 400–600 kpc. Previous temperature profiles for Hydra A were derived from *ASCA* data by Ikebe et al. (1997), Markevitch et al. (1998), and White (2000). The primary difference between the *Chandra* and *ASCA* profiles is the lower gas temperature in the center of the cluster and the steeper temperature gradient. This is to be expected given the greater spatial resolution of *Chandra*.

3.2. Abundance Gradients

Intermediate temperature clusters like Hydra A emit strongly in both Fe L and Fe K lines. To determine the Fe abundance of the gas, we fitted the ACIS-S spectra in a broad energy band from 0.5–7.5 keV, that includes emission from both Fe L and K lines, and a hard band from 2.0–7.5 keV, that includes only Fe K lines. The best-fit Fe abundances are shown in Figure 10 based on an absorbed VMEKAL model with N_{H} , all α -processed elements, and Fe treated as free parameters. This figure shows that there is excellent agreement between these two determinations and that there is a strong Fe abundance gradient within the central 100 kpc. The agreement between the two Fe estimates shows that potential low-energy complications, e.g., excess absorption or cooling gas, do not have a significant impact on the Fe abundance measurements. An Fe abun-

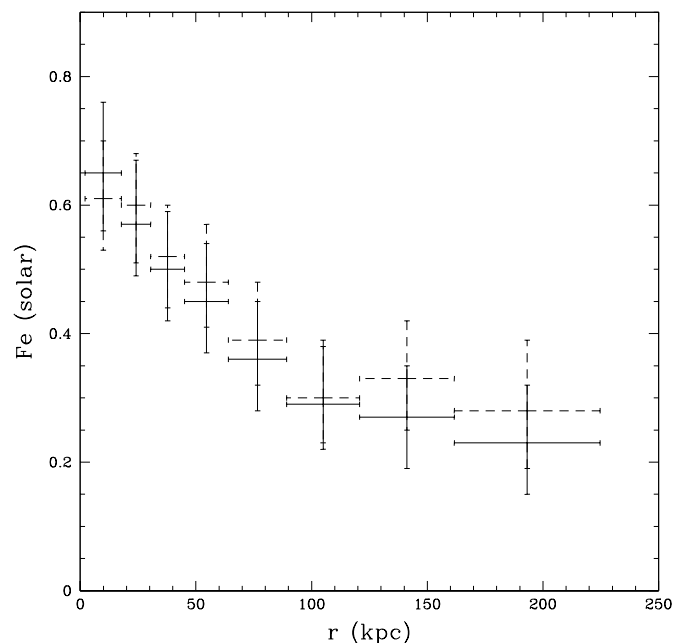


FIG. 10.—Fe abundance profile relative to the solar value (4.68×10^{-5} by number relative to H) for the Hydra A cluster derived from the ACIS-S observation. The solid lines show the best fit to a single-temperature model in the 0.5–7.5 keV band, while the dashed lines show the best fit to a single-temperature model in the 2.0–7.5 keV band. Error bars are shown at the 90% confidence level for one interesting parameter.

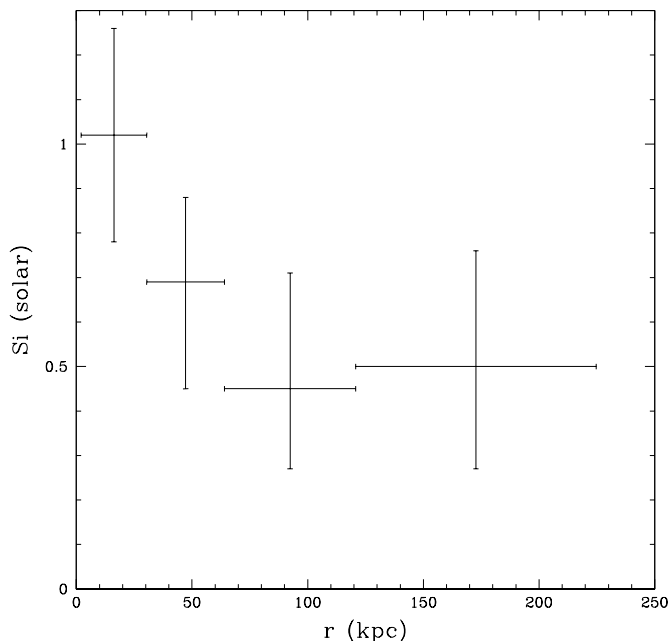


FIG. 11.—Si abundance profile relative to the solar value (3.55×10^{-5} by number relative to H) for the Hydra A cluster derived from the ACIS-S observation. Error bars are shown at the 90% confidence level for one interesting parameter.

dance gradient in Hydra A also was reported based on *ASCA* data by Ikebe et al. (1997) and Finoguenov et al. (2000). While the magnitude of the measured Fe gradient is comparable between the *ASCA* and *Chandra* analysis, only *Chandra* has the spatial resolution to resolve the Fe distribution within the central few arcminutes. Figure 10 shows that the Fe gradient is confined to the central 100 kpc, while analysis of the *ASCA* data implied a flatter gradient extending to 200–300 kpc.

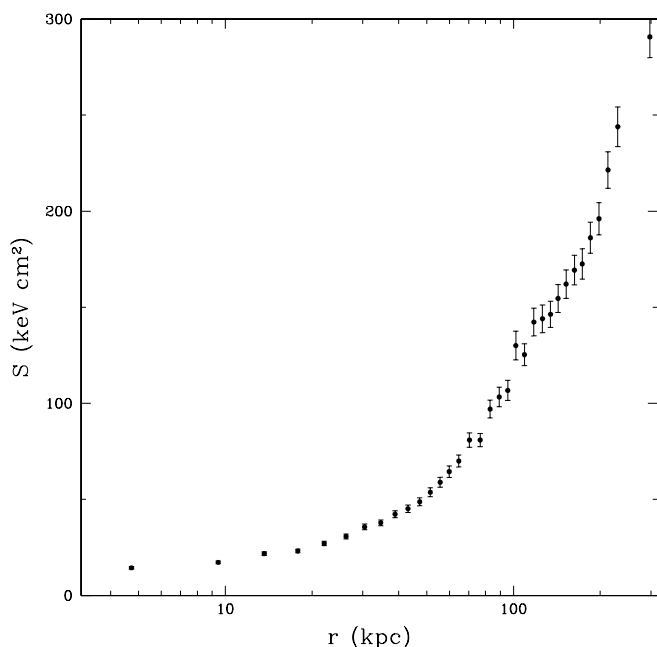


FIG. 12.—Gas entropy ($S = kT/n_e^{2/3}$) profile derived from the deprojected density and power-law temperature profiles.

Significant constraints on the Si abundance can be obtained only from the set of spectra with 40,000 net counts. Figure 11 shows that the Si distribution is similar to the Fe distribution with a rapid decline within the central 100 kpc. Beyond 100 kpc, the Si abundance is essentially constant within the uncertainties. Because of the large errors on the Si abundance, it is difficult to make conclusive comments on the radial variation in the Si/Fe abundance ratio.

3.3. Deprojected Temperature and Entropy Profiles

The projected temperature profile shows that there is a strong positive gradient within the central 200 kpc (see Fig. 9). To determine the deprojected temperature profile in this region, we first fitted the outer ACIS-S spectrum to a single MEKAL model, with N_H and the abundance of heavy elements treated as free parameters. We then fitted the spectrum extracted from the next smallest annulus to a two-temperature model, with the parameters of one of the models frozen to the best-fit values derived from the outermost spectrum. The normalization of this model is adjusted to account for the volume within the outer shell projected along the line of sight toward the next smallest shell. This procedure is continued inward, adding a MEKAL model for each successive spectrum, with the free parameters of the remaining models frozen to the previously determined best-fit values. Finally, the best-fit temperatures, along with their associated errors, are fitted to a power-law profile given by $T(r) = T_0(r/10 \text{ kpc})^p$. The best fit is obtained with $T_0 = 2.73 \pm 0.07 \text{ keV}$ and $p = 0.124 \pm 0.014$. The gas entropy is determined from the deprojected density profile and the best-fit power-law temperature profile (see Fig. 12). This figure shows that the gas is approximately isentropic within the central 30 kpc, and the gas entropy only increases significantly at larger radii.

4. MASS DISTRIBUTIONS

The gravitating mass distribution is determined from the equation of hydrostatic equilibrium, given by

$$M_{\text{tot}}(< r) = -\frac{kTr}{\mu m_p G} \left(\frac{d \ln \rho_{\text{gas}}}{d \ln r} + \frac{d \ln T}{d \ln r} \right). \quad (1)$$

We use a parametric form for the gas temperature (the best power-law profile given above) along with the non-parametric density profile shown in Figure 3 to determine the mass distribution in Hydra A. To reduce the noise in the $d \ln \rho_{\text{gas}}/d \ln r$ term, we determine the gradient at r_j by differencing the densities at r_{j-2} and r_{j+2} . The density gradients are assumed to be equal in the innermost three points and the outermost three points. The cumulative gravitating mass, gas mass, Fe mass, and Si mass are shown in Figure 13. In addition, we show the gravitational potential in Figure 14, with the potential normalized to zero at 250 kpc. Of the four mass components shown in Figure 13, the Fe and Si are the most centrally concentrated, the gas is the most extended, and the dark matter is intermediate between the two. The central concentration of a given mass component can be quantized by calculating the ratio of the mass within 100 kpc to that within 200 kpc. For the gas and dark matter this ratio is 27% and 33%, respectively. Also shown in Figure 13 are the integrated Fe and Si masses that would be derived under the assumption of a uniform distribution of metals using the best-fit values derived from a spectral analysis of the integrated emission within the central 200

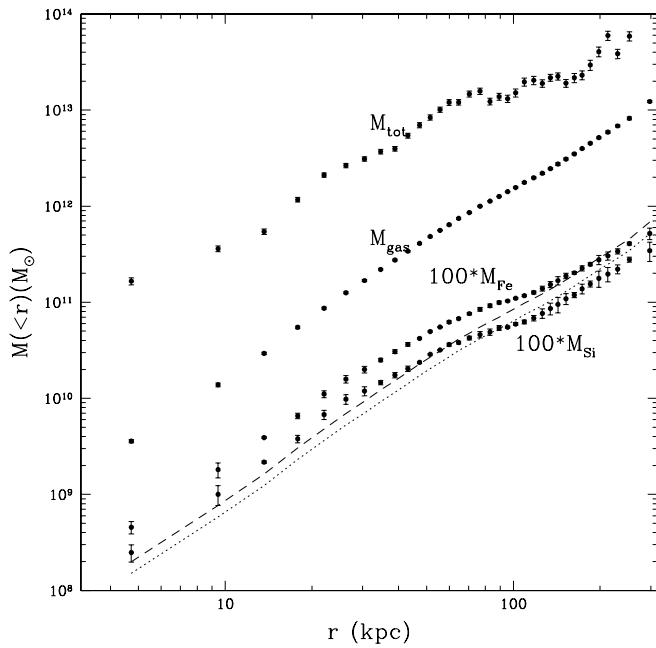


FIG. 13.—Integrated total gravitating mass, gas mass, Fe mass, and Si mass distributions. The Fe and Si masses are shown at 100 times their true values. The dashed and dotted lines are the Fe and Si masses that would be inferred assuming a uniform distribution of heavy elements.

kpc. This figure shows that the assumption of a uniform distribution of heavy elements underestimates the Fe and Si masses by a factor of 2 at small radii and overestimates the Fe and Si masses by approximately 30–50% at large radii. This has important consequences concerning the chemical enrichment history of clusters.

The variation in the gas mass fraction with radius is shown in Figure 15. This figure shows that the gas mass fraction increases from 4% at 10 kpc (several times greater

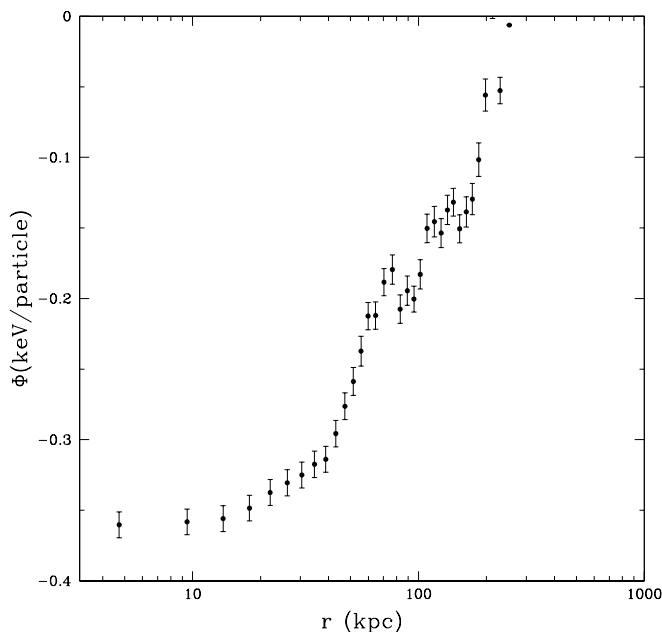


FIG. 14.—Gravitational potential in the Hydra A cluster derived from the total mass distribution shown in Figure 8. The potential is normalized to zero at 250 kpc.

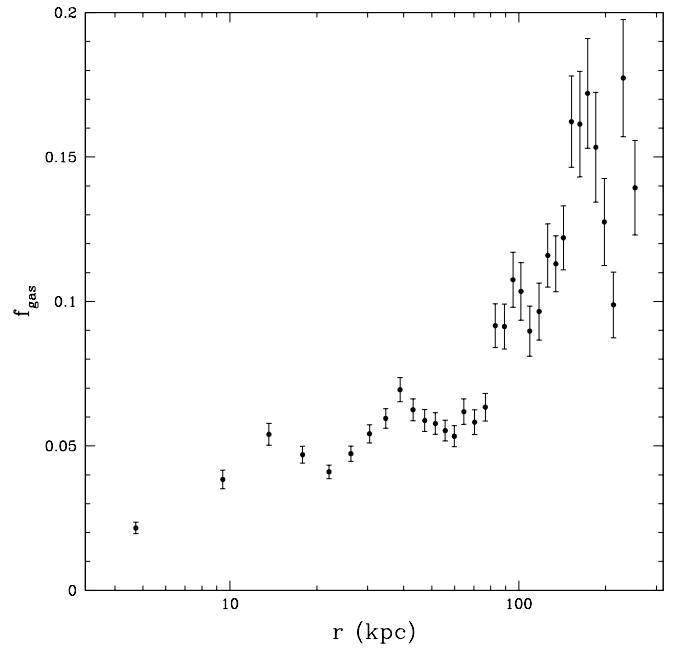


FIG. 15.—Gas mass fraction vs. radius in the Hydra A cluster

than a typical isolated giant elliptical) up to 15% at 200 kpc, which is typical of rich clusters (David, Jones, & Forman 1995; White & Fabian 1995; Evrard 1997). The effects of the cluster environment on the X-ray characteristics of the central dominant galaxy will be discussed in § 6. The blue light within the central 47 kpc of Hydra A is $L_B = 9.2 \times 10^{10} L_\odot$ (see § 6), which gives $M/L_B = 76 M_\odot/L_\odot$. Assuming $M/L_B = 8 M_\odot/L_\odot$ for the stellar population of Hydra A gives a stellar mass fraction of 10% for a total baryon fraction of 15%. The total baryon fraction within the central 47 kpc is thus comparable to the gas mass fraction beyond 100 kpc.

4.1. Comparison with Theoretical Expectations

There has been a great deal of discussion in the literature concerning the expected density profile of CDM halos and the observed density distribution of galaxies and clusters. Based on an extensive set of numerical simulations of the formation and evolution of dark matter halos, Navarro, Frenk, & White (1995, 1997) concluded that dark matter halos have a universal density profile given by

$$\frac{\rho_d}{\rho_{\text{crit}}(z)} = \frac{\delta_c}{(r/r_s)^\gamma [1 + (r/r_s)^\alpha]^{(\nu-\gamma)/\alpha}}, \quad (2)$$

where $\alpha = 1$, $\nu = 3$, and $\gamma = 1$, $\rho_{\text{crit}}(z) = 3H^2/8\pi G$ is the critical density of the universe at the cluster redshift; r_s is a characteristic radius; and δ_c is the central overdensity, which can be expressed in terms of the concentration parameter ($c = r_{200}/r_s$, where r_{200} is the radius within which the mean halo density is $200\rho_{\text{crit}}$) as

$$\delta_c = \frac{200}{3} \frac{c^3}{[\ln(1+c) - c/(1+c)]}. \quad (3)$$

The NFW density profile varies from $\rho_d \propto r^{-1}$ at small radii to $\rho_d \propto r^{-3}$ at large radii. For a given cosmology, the concentration parameter decreases with increasing halo mass. Higher resolution simulations by Moore et al. (1999)

found a density profile similar to the NFW profile at large radii, but with a steeper density profile at small radii ($\rho_d \propto r^{-1.5}$) corresponding to $\alpha = 1.5$, $\nu = 3$, and $\gamma = 1.5$. Jing (2000) examined how the density profile depends on the dynamic state of dark matter halos. He found significant variations in the density profile of halos at the present epoch in his simulations. By fitting the simulated halo density profiles with an NFW profile, he found that the observed distribution in the concentration parameter can be described by a lognormal function. Virialized systems have small residuals when fitted with the NFW profile and a mean concentration parameter comparable to that predicted by NFW. Unrelaxed systems produce larger residuals and smaller concentration parameters when fitted to the NFW profile.

The primary difficulty encountered when these theoretical predictions for CDM halos are compared to observations is the inferred presence of flat cores in the density profiles of dwarf and low surface brightness galaxies (Flores & Primack 1994; Burkert 1995; Moore et al. 1999). For example, Moore et al. showed that the H I rotation curves of low surface brightness galaxies are best fitted with a modified King profile ($\alpha = 2$, $\nu = 3$, and $\gamma = 0$). However, van den Bosch et al. (2000) argue that the resolution of the present H I data is insufficient to decouple the disk and halo components of low surface brightness galaxies and cannot be used to place significant constraints on the central density distribution. They suggest that flat cores may only exist in low-mass dwarf systems, which are more susceptible to the effects of supernovae-driven winds. A recent paper by Swaters, Madore, & Trewella (2000) demonstrates that high-resolution H α rotation curves of low surface brightness galaxies do indeed increase faster with radius than lower resolution H I rotation curves. On cluster scales, Tyson, Kochanski, & Dell'Antonio (1998) analyzed strong lensing data on CL 0024+1654 and found evidence for a flat core, while Broadhurst et al. (2000) concluded that a flat core is not required by the strong lensing data for this cluster. However, Shapiro, & Iliev (2000) show that the mass distribution used by Broadhurst et al. predicts a galaxy velocity dispersion much greater than that observed, while the mass distribution derived by Tyson et al. is consistent with the observed velocity dispersion.

The cumulative gravitating mass distribution of the Hydra A cluster in Figure 13 shows that there is little evidence for a flat core with a core radius greater than approximately 30 kpc. We cannot exclude the presence of a smaller core, since we are only able to constrain the gas temperature within two annuli in this region (see Fig. 9). Between 30 and 200 kpc, the integrated mass scales as $M(<r) \propto r^{1.7}$, or equivalently, $\rho_d \propto r^{-1.3}$, which is slightly steeper than the NFW profile and slightly flatter than the Moore profile. The mass profile is actually in very good agreement with simulations that incorporate the effects of gas cooling and star formation by Lewis et al. (2000), who find $\rho_d \propto r^{-1.4}$. Fitting the mass distribution in Hydra A to these analytic models we obtain best-fit parameters of $\delta_c = (7.48 \pm 0.14) \times 10^4$ and $r_s = 77 \pm 10$ kpc, $\delta_c = (4.02 \pm 0.36) \times 10^3$ and $r_s = 234 \pm 25$ kpc, and $\delta_c = 4.54 \times 10^5$ and $r_s = 26 \pm 2$ kpc, for the NFW, Moore, and modified King profiles, respectively. The best-fit δ_c uniformly decreases with the level of mass concentration in the models. Equation (3) is strictly valid only for the NFW profile and yields a concentration parameter of $c_{\text{NFW}} = 12.3 \pm 0.18$. The best-fit pro-

files for the three mass models are shown along with the cumulative gravitating mass in Figure 16. This figure shows that both the NFW and Moore profiles provide reasonable fits to the data, supporting the results of numerical CDM simulations. Based on the rotation curve of the H α disk in Hydra A published in Melnick et al. (1997), the gravitating mass within 3.7 kpc is $3.9\text{--}7.8 \times 10^{10} M_\odot$, for inclination angles ranging from 0° to 45° . The addition of this point in Figure 16 favors the NFW profile. The gravitational potential of the central galaxy should not affect our results since the gas temperature at 10 kpc is 2.7 keV. This is significantly hotter than the virial temperature of the central galaxy ($kT \approx 0.6$ keV) derived from the stellar velocity dispersion of 307 km s^{-1} (McElroy 1995).

The Hydra A cluster is approximately a $10 M_*$ cluster. For the CDMA cosmology used by NFW (essentially the same as that used to estimate distance dependent quantities in this paper) the predicted concentration parameter for a $10 M_*$ halo is 4.0, which is only one-third of the observed value. The high observed concentration of the Hydra A cluster implies an improbable collapse redshift of $z \approx 4$ (using the expression given in the appendix of NFW). The high concentration parameter in Hydra A cannot be resolved with NFW by appealing to the possible presence of residual substructure since Jing (2000) finds that unrelaxed clusters yield concentration parameters lower than the NFW values. Based on the NFW simulations, a concentration parameter of $c_{\text{NFW}} \approx 12$ is expected for a $0.01 M_*$ dark matter halo, which actually corresponds to the integrated mass within the central 10 kpc of Hydra A. In summary, we find that the total density distribution in the central region of the Hydra A cluster is fairly steep with $\rho_d \propto r^{-1.3}$. The observed mass distribution is in reasonable agreement with CDM simulations, but the derived concentration parameter is more representative of a galaxy-size halo rather than a

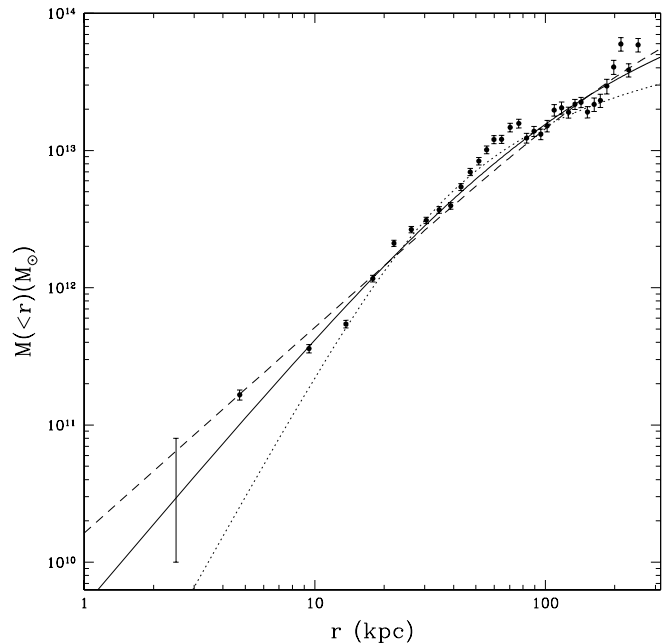


FIG. 16.—Comparison of the best-fit NFW profile (solid line), Moore profile (dashed line), and modified King profile (dotted line) to the cumulative gravitating mass distribution in Hydra A. The point at 3.7 kpc is the mass inferred from the rotation curve of the H α disk (Melnick et al. 1997).

rich cluster of galaxies. This analysis will be extended to a sample of rich clusters in Robertson & David (2001, in preparation).

5. EVIDENCE FOR FEEDBACK IN THE COOLING FLOW

The morphological mass accretion rate across a spherical shell j can be estimated from

$$\dot{M}_j = \frac{L_j}{(\Delta H_j + \Delta \phi_j)}, \quad (4)$$

where L_j is the deprojected X-ray luminosity within shell j , and ΔH_j and $\Delta \phi_j$ are the changes in gas enthalpy and gravitational potential across shell. Equation (4) gives the mass accretion rate assuming steady state and no mass deposition, in agreement with the spectroscopic results beyond 30 kpc. This equation differs from that used in, e.g., Allen (2000) because we do not include the X-ray emission arising from gas condensing out of the flow. Figure 17 shows the resulting morphological \dot{M} profile using the deprojected X-ray luminosity, the best-fit power-law temperature profile, and the gravitational potential shown in Figure 14. Equation (4) is only strictly valid beyond 30 kpc and the only self-consistent solution is one with a constant \dot{M} , which is in fact the derived solution. Beyond 30 kpc, the morphological \dot{M} is nearly constant at $300 M_\odot \text{ yr}^{-1}$ (see Fig. 17). This result provides strong evidence that there is a steady, (nearly) homogeneous cooling flow beyond 30 kpc. Inside 30 kpc, there is spectroscopic evidence for multiphase gas, but the spectroscopic mass deposition rate is more than a factor of 10 less than the morphological mass accretion rate at 30 kpc. Figure 17 also shows that the \dot{M} calculation is not very sensitive to the details of the gravitational potential. The change in gas enthalpy within the central 200 kpc is 2.5 keV per particle, while the change in gravitational potential is only 0.5 keV per particle (see Fig. 14).

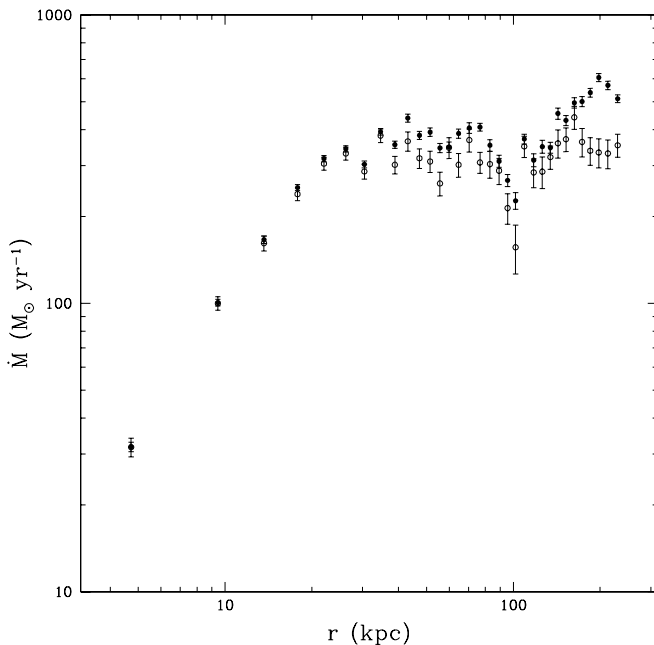


FIG. 17.—Morphological \dot{M} profile derived from the deprojected surface brightness and temperature profiles. Filled circles only include the change in gas enthalpy across a shell, while the open circles include the change in enthalpy and gravitational potential in eq. (4).

The discrepancy between the morphological mass accretion rate at 30 kpc and the spectroscopic \dot{M} within this radius suggests that some heating mechanism may be suppressing the accretion of the cooling gas. Many heating mechanisms have been proposed over the past two decades in the hopes of significantly reducing the morphological mass accretion rate of cooling flow clusters. One of the most promising is heating from a central radio source, since approximately 70% of cooling flow clusters have active central galaxies. Heinz, Reynolds, & Begelman (1998) showed that shock heating by expanding relativistic plasma can induce observable X-ray features in the ambient cluster gas. Tucker & David (1997) showed that collective heating effects from relativistic plasma can generate a duty cycle between accretion and outflow that can significantly reduce the integrated mass accreted over the lifetime of a cluster. A similar model based on Compton heating by a central active galactic nucleus (AGN) in ellipticals has been developed by Ciotti & Ostriker (1999). Recently, Soker et al. (2000) proposed a moderate cooling flow model due to heating by radio sources that primarily affects the cooling rate of gas in the outer regions of cooling flows, significantly reducing their time-averaged mass accretion rates. In addition to direct heating, Churazov et al. (2000) have shown that hot buoyant bubbles produced through shock heating by the expanding central radio source can dredge up cold material from the center of the cluster. In this paper we further examine the role of weak and strong shock heating by a central radio source and discuss the development of a convectively unstable region in cooling flows that can significantly reduce the net inflow of cooling gas.

5.1. Heating by Weak Shocks

The lack of evidence for gas deposition and the flattening of the entropy profile near the center of the cluster suggest that the cluster gas is being heated appreciably by the radio source. On the other hand, if this is the case, then the temperature decline toward the center of the cluster is surprising, since it is an indication of cooling. If the age of this cluster since its last major merger is about 10^{10} yr, then, in order to prevent the deposition of $300 M_\odot \text{ yr}^{-1}$ of cooled gas, the heating power needs to make up for the radiative heat loss within the central 150 kpc of the cluster, i.e., about $2.5 \times 10^{44} \text{ ergs s}^{-1}$. This is plausible, since it is only 4 times the current radio power (Ekers & Simkin 1983). Here we consider some of the details of how heating by shocks from the radio source can prevent gas deposition.

It is generally argued that the mechanical power of expanding radio jets exceeds the radio power by a substantial factor. This leads to models in which the expanding radio jets drive a roughly spherical shock into the surrounding gas (Heinz et al. 1998). Shock heating due to such radio outbursts is a good candidate for the required heating. The typical lifetime of cluster center radio sources is 10^7 – 10^8 yr (e.g., Taylor et al. 1990), so that the high incidence of central radio sources in cooling flow clusters requires that they undergo frequent repeated outbursts. We begin by considering the accumulated heating effect of repeated weak shocks due to these outbursts.

The shock jump conditions may be written as

$$\frac{p_1}{p_0} = 1 + \frac{2\gamma}{\gamma + 1} y \quad (5a)$$

and

$$\frac{\rho_1}{\rho_0} = \frac{(\gamma + 1)(1 + y)}{\gamma + 1 + (\gamma - 1)y}, \quad (5b)$$

where p is the pressure, ρ the density, γ the ratio of specific heats, and subscripts 0 and 1 refer to preshock and postshock conditions, respectively. The quantity $y = \rho_0 v_s^2 / (\gamma p_0) - 1$ is the square of the shock Mach number minus 1, so that $y \ll 1$ for weak shocks and $y \gg 1$ for strong shocks. The specific entropy jump for a weak shock is then

$$\Delta S \simeq \frac{2\gamma k}{3(\gamma + 1)^2 \mu m_H} y^3. \quad (6)$$

For a weak shock the kinetic energy of the shocked gas is negligible compared to its increase in thermal energy. Assuming that the postshock pressure excess is roughly uniform, we can therefore estimate the shock strength by equating the mechanical energy that drives the shock to the excess thermal energy within the volume V encompassed by the shock, i.e.,

$$E_T \simeq \frac{1}{\gamma - 1} (p_1 - p_0) V. \quad (7)$$

Using the shock jump condition for p_1 , this gives an expression for y , which determines the specific entropy jump in the shock:

$$\Delta S \simeq \frac{(\gamma - 1)^3 (\gamma + 1) k}{12\gamma^2 \mu m_H} \left(\frac{E_T}{p_0 V} \right)^3. \quad (8)$$

The requirement that the shock is weak is $E_T \ll p_0 V$. If the average repetition frequency of outbursts is $1/\tau$, then the average heating rate per unit volume due to weak shocks is

$$H = \rho_0 T_0 \frac{\Delta S}{\tau} = \frac{(\gamma - 1)^3 (\gamma + 1) p_0}{12\gamma^2 \tau} \left(\frac{E_T}{p_0 V} \right)^3. \quad (9)$$

Comparing this to the radiative cooling rate per unit volume, $R = \rho_0^2 \Lambda(T_0)$, gives

$$\frac{H}{R} = \frac{(\gamma - 1)^4 (\gamma + 1)}{12\gamma^2} \left(\frac{t_{c,0}}{\tau} \right) \left(\frac{E_T}{p_0 V} \right)^3, \quad (10)$$

where the cooling time is $t_{c,0} = p_0 / [(\gamma - 1)R]$.

Taking $\Lambda(T) \propto T^\alpha$, the ratio $H/R \propto \rho_0^{-4} T_0^{-(2+\alpha)} r^{-9}$, so that, if the gas is nearly isothermal, H/R increases with r only if ρ_0 decreases more rapidly than $r^{-9/4}$. This is significantly steeper than the observed r dependence of ρ in the region of interest, so that H/R is a decreasing function of r within the central 150 kpc of the Hydra A cluster. Thus, heating by weak shocks can only dominate cooling at small r . The shock heating rate would dominate the cooling rate everywhere inside the point where they balance, i.e., $r = r_b$, where $H = R$. At small r the net heating time, $p_0 / [(H - R)] \simeq p_0 / [(\gamma - 1)H]$, would then be shorter than the cooling time and the gas is heated rapidly. Where the net heating is substantial it leads to convective instability because the heating time is shortest at small r . Where the heating is not too fast (which includes the region where the shocks are weak), convective instability drives mixing that will result in a nearly isentropic core.

Details of the process leading to convection are complex, since as the gas is heated it is redistributed and the pressure

and density profiles are flattened in order to maintain hydrostatic equilibrium. At r_b the entropy of the gas remains fixed. To get a rough idea of what part of the gas becomes isentropic we use the crude estimate

$$t_S = \frac{S(r_b) - S(r)}{|dS/dt|_r} \quad (11)$$

for the time it takes for the entropy of the gas at r to rise to the entropy of gas at r_b . Assuming that the gas is (initially) isothermal and $\rho_0(r) \propto r^{-\beta}$, the results above then give

$$t_S = \frac{(\gamma - 1)\beta \ln(r_b/r)}{(r_b/r)^{9-4\beta} - 1} t_{c,0}(r). \quad (12)$$

Taking $\beta = 1.5$, roughly appropriate for the region outside the isentropic core in Hydra, this gives $t_S \simeq 0.1 t_{c,0}$ at $r = r_b/2$. Since the cooling time is already relatively short in this region, this suggests that the bulk of the gas within r_b should be isentropic.

5.2. Heating by Strong Shocks

For the purpose of this discussion, heating by strong shocks adds one significant complication, that a single strong shock can cause a substantial entropy rise. This can produce ‘‘bubbles’’ of gas with such high entropy that they rise well outside the isentropic core region (r_b) before arriving at their equilibrium position. Convection still creates an isentropic core within a few free-fall times after passage of the shock, but some gas may rise well outside this core.

We identify the roughly isentropic region at $r \lesssim 30$ kpc as that where the average shock heating rate exceeds radiative cooling. This determines the strength of shocks due to the radio outbursts as follows. The shock repetition time, τ , is highly uncertain, but, as mentioned above, it needs to be comparable to the radio-source lifetime to account for the high incidence of central radio sources in cooling flow clusters. Taking $\tau = 10^8$ yr, the cooling time at $r = 30$ kpc is too short relative to the shock repetition time for weak shocks to balance the radiative cooling rate. Instead we must appeal to stronger shocks to provide the heating, but in that case balancing the heating rate against the average cooling rate is complicated by the large excursions in gas properties caused by the shocks. Provided that the gas is reasonably close to hydrostatic equilibrium most of the time, we can use the ratio of the average rate of change of entropy due to shocks to that due to radiative heat loss, i.e.,

$$Q = \frac{\Delta S/\tau}{R/(\rho T)} = \frac{t_c}{\tau} \left[\ln \left(1 + \frac{2\gamma y}{\gamma + 1} \right) + \gamma \ln \left\{ \frac{\gamma + 1 + (\gamma - 1)y}{(\gamma + 1)(1 + y)} \right\} \right]. \quad (13)$$

This is essentially the same as the expression for H/R above, except that here t_c is an ‘‘average’’ cooling time, which we assume to be close to the observed cooling time, and the expression for the entropy jump in terms of y is exact. Taking $t_c/\tau \simeq 10$ at $r = 30$ kpc and $\gamma = 5/3$ then requires $y \simeq 2$ to make $Q = 1$ there. If the observed pressure is close to the preshock pressure, p_0 , equation (7) still gives a reasonable estimate for the shock energy. Taking $y = 2$, $kT = 3.3$ keV and $n_e = 0.027$ cm $^{-3}$, as observed at $r = 30$ kpc, we get $E_T \simeq 3.4 \times 10^{60}$ ergs. For $\tau = 10^8$ yr, this makes the average shock power about 1.1×10^{45} ergs s $^{-1}$. This is

to be compared with the total power radiated by the gas within $r = 150$ kpc, i.e., about 2.5×10^{44} ergs s^{-1} . Given the large uncertainty in the shock energy, the agreement is close, suggesting that the radiative cooling rate in the central 150 kpc and shock heating due to radio outbursts are tied together. Note that these numbers agree for $\tau \simeq 4 \times 10^8$ yr.

Direct shock heating does not fully explain the lack of gas deposition, since the argument above suggests that it is insignificant compared to radiative cooling outside the central 30 kpc. In that case, with no heat source to make up for radiative losses, cooling gas from the region between 30 and 150 kpc flows into the central 30 kpc at about $300 M_{\odot} y^{-1}$. If the system is to remain in an approximately steady state, then an equivalent amount of gas must be driven back into the region surrounding the core by each radio outburst. Again, taking $\tau \simeq 10^8$ yr, this means that each radio outburst must drive about $3 \times 10^{10} M_{\odot}$ of gas from the core. This is roughly 30% of the mass of gas in the 30 kpc core. The core gas is roughly isentropic, so a strong shock can certainly cause a substantial entropy inversion, i.e., gas that can rise well outside the core. It is not clear that as much as 30% of the gas can be lifted from the core by a single shock. More accurate numbers and more detailed modeling are needed to determine if this is possible.

The detectability of the hot bubbles rising from the core is determined by their size and speed. They are unstable, and will tend to break up and mix with ambient gas as they rise. Given that they are unlikely to rise faster than the thermal speed, i.e., about 730 km s^{-1} , they would take at least 1.3×10^8 yr to reach 100 kpc. This means that some hot bubbles should be present most of the time during the cycles between outbursts. The mass of this hot gas is a few percent of the total gas mass within 100 kpc, so it may be difficult to detect when it is well away from the core. On the other hand, it would be easily detected after the passage of the shock, while it still forms a coherent, overpressured body. Because the hot gas rises relatively quickly (speeds of order the sound speed), it fills a relatively small volume compared to inflowing cooling gas (which is highly subsonic). Thus, apart from some mixing, the outflowing hot bubbles may have relatively little impact on the steady inflow due to cooling.

The shock itself would be easily detectable anywhere within the 30 kpc core region (for $y = 2$, the density jump is a factor of 2 and the temperature jump a factor of 1.75). At $r = 30$ kpc, the shock speed is $v_s = (3\gamma kT_0/\mu m_H)^{1/2} \simeq 1560 \text{ km s}^{-1}$, so that the shock takes about 10^7 yr to reach $r = 30$ kpc. This means that the shock is easily visible for at least 10% of the time, or in at least 10% of all comparable sources. If this model applies to many clusters, such shocks should be detected soon. If shocks are not detected, then shocks from the radio source cannot be responsible for preventing the deposition of cooled gas and we would need to find another heat source.

This description of gas flows around Hydra A presumes that the system remains in a quasi-steady state. Using the relationship between stellar velocity dispersion and the mass of the central black hole (Gebhardt et al. 2000; Ferrarese & Merrit 2000), the central galaxy in Hydra A should have a nuclear black hole with a mass of $\simeq 5 \times 10^8 M_{\odot}$. As cooling gas accumulates in the cluster core, the density of gas around this black hole increases, presumably driving up the accretion rate until there is a radio outburst. The out-

burst heats and drives away the central gas, thereby reducing the accretion rate and quenching the outburst. This provides a feedback mechanism connecting the average mechanical power of radio outbursts to the total radiative power within the region where there would otherwise be a cooling flow (i.e., where the cooling time is shorter than the age of the system). The critical condition is that the net mass flow rate into the core region is kept small (the spectroscopically determined mass deposition rate is less than 10% of the rate at which cooling gas is delivered to the core in this system). In the absence of a detailed model for the outflow, it is not clear that this requires very close balance between the average mechanical heating rate and radiative cooling. The mechanical heating rate must at least match the cooling rate to prevent accumulation of gas in the core, but any energy deposited outside the cooling flow region has little impact on the current cooling rate and so is not closely controlled by this feedback process. Thus the mechanical heating rate may exceed the radiative heat loss in the cooling flow region. In that case, the net heat input to the cluster could be significant for the cluster as a whole. For example, an average excess heating rate of 10^{45} ergs s^{-1} for 10^{10} yr deposited in $10^{14} M_{\odot}$ of gas amounts to 1.0 keV per particle. This is close to what is required to account for “preheating” in clusters (e.g., Wu, Fabian, & Nulsen 2000), and the excess heating rate per unit mass could well be greater at early times, when AGNs were more active.

In summary, for Hydra A, gas in the region from about 30 to 150 kpc is cooling steadily and flowing inward, essentially as a (homogeneous) cooling flow at a rate of about $300 M_{\odot} \text{ yr}^{-1}$. Repeated outbursts from the central radio source drive shocks into the cluster and heat the region $r < 30$ kpc strongly enough to prevent all but about 10% of the inflowing gas from cooling to low temperatures. Because it is heated from the center, the 30 kpc core is convective and remains roughly isentropic. During each outburst, gas near the center of the core is heated so much that it rises well outside the core before finding its equilibrium position. On average, the resulting mass outflow nearly balances the cooling inflow to the core. The rate of accretion onto the central black hole that is the source of the radio outbursts is determined by the state of gas in the core. This link provides the feedback that keeps the net gas flow rate into the core close to zero. This model has much in common with the model of Tabor & Binney (1993) for gas flows in elliptical galaxies. Shock heating fits well with current models of radio outbursts, but if shocks are the main heat source in the core, they soon should be detected in some systems. If shocks are not a major heat source, then any other process by which the AGN heats the cluster from its center would produce a similar general picture. The major factor distinguishing such models is the extent to which the heating is intermittent. Continuous heating would tend to make the whole of the region, which would otherwise be depositing cool gas, isentropic.

6. SOURCE OF THE CENTRAL ABUNDANCE GRADIENT

Between 100 and 200 kpc the Fe and Si abundances are essentially constant with a Si/Fe abundance ratio twice solar (see Figs. 10 and 11). This agrees with earlier results from *ASCA* observations of clusters indicating a significant role of Type II supernovae (SNe II) in the global enrichment of cluster gas (Mushotzky et al. 1996; Fukazawa et al. 1998). Within the central 100 kpc there is an excess of Si and Fe

above the globally averaged values. Using an average Si abundance of 0.5 solar gives an excess Si mass of $\Delta M_{\text{Si}} \simeq 1.4 \times 10^8 M_{\odot}$ within the central 100 kpc. Using an average Fe abundance of 0.25 solar gives an excess Fe mass of $\Delta M_{\text{Fe}} \simeq 3.0 \times 10^8 M_{\odot}$ within the same region. The ratio $\Delta M_{\text{Si}}/\Delta M_{\text{Fe}} = 0.5$ is essentially consistent with the solar ratio of 0.38. Fukazawa et al. (2000) and Irwin & Bregman (2001) find that central metal excesses are common in cD clusters with cooling flows based on *ASCA* and *BeppoSAX* observations. Fukazawa et al. also find that the excess metals have a nearly solar abundance ratio. However, *ASCA* was unable to resolve the distribution of excess metals in clusters and accurately determine the Fe and Si masses.

The primary sources of metals in the central region of the Hydra A cluster are stellar winds and Type Ia Supernovae (SNe Ia) from the old stellar population in the central galaxy and SNe II from the young stellar population in the blue disk reported by McNamara (1995). A central abundance excess will develop even if the gas is static due to the more extended distribution of the gas compared to the light from the central galaxy. Based on the spectroscopic mass deposition rate, only 10% of the gas within the central 100 kpc should have condensed out of the hot gas over the cluster lifetime with the remainder being continuously enriched. Convection will flatten the abundance gradients to some extent but will not completely erase all gradients. Convection will also dredge up material from within the central 10 kpc in the central galaxy that should be enriched by stars with supersolar abundances, typical of the central regions of elliptical galaxies.

To determine the origin of the excess metals, we first compute the number of supernovae required to produce the excess Si and Fe masses. We adopt the supernovae yields given in Finoguenov et al. (2000), that are averaged over a Salpeter IMF and given by $y_{\text{Fe}}(\text{Ia}) = 0.74 M_{\odot}$, $y_{\text{Fe}}(\text{II}) = 0.070 M_{\odot}$, $y_{\text{Si}}(\text{Ia}) = 0.158 M_{\odot}$, and $y_{\text{Si}}(\text{II}) = 0.133 M_{\odot}$. The excess Fe mass requires 4×10^8 SNe Ia. Since the Si yield is nearly independent of supernova type, we can simply assume a yield of $0.14 M_{\odot}$ per supernovae. To produce the excess Si mass requires 10^9 total supernovae, or 6×10^8 SNe II. Assuming 1 SNe II for every $100 M_{\odot}$ of star formation (appropriate for a Salpeter IMF) requires an average star formation rate of $60 M_{\odot} \text{ yr}^{-1}$ over the past 10^9 yr. McNamara (1995) showed that the observed luminosity and colors of the blue disk in Hydra A could only be reproduced by continuous star formation over the past 10^9 yr with star formation rates of $\lesssim 1 M_{\odot} \text{ yr}^{-1}$. Higher star formation rates are also permitted by the data, but only when coupled with shorter periods of star formation. It is thus unlikely that the central metal excess in Hydra A was produced by star formation in the blue disk.

The total light of the central galaxy can be estimated from the photometric observations by Peterson (1986) who gives $M_V = -23.65$ and $B-V = 1.08$ within a 1.65 diameter aperture. This gives $L_B = 9.2 \times 10^{10} L_{\odot}$ within the central 47 kpc radius using our adopted cosmology. To reproduce the observed excess Fe mass requires an average SNe Ia rate of $4 \times 10^{-13} L_B^{-1} \text{ yr}^{-1}$ over the past 10^{10} years, which is approximately twice Tammann's (1974) rate of $2.2 \times 10^{-13} L_B^{-1} \text{ yr}^{-1}$. Tammann's rate is usually considered to be a factor of 4 too high (van den Bergh, McClure, & Evans 1987), but our estimate does not include the contribution from stars beyond the central 50 kpc. Given the large uncer-

tainty in the present epoch SNe Ia rate in elliptical galaxies, the excess Fe can be accounted for by SNe Ia from the old stellar population of the central galaxy.

Mass loss from stars as they evolve off the main sequence also can be a significant source of metals. Using the Faber & Gallagher (1976) mass loss rate of $10^{-11} M_{\odot} L_V^{-1} \text{ yr}^{-1}$ gives a stellar mass loss rate of $2.5 M_{\odot} \text{ yr}^{-1}$ within the central 47 kpc. If this gas has a solar abundance of heavy elements, and the present mass loss rate is assumed to hold over the lifetime of the galaxy, then stellar mass loss can only produce about 15% of the observed Si and Fe excess, and only about 6% of the total gas mass within this radius.

There is considerable evidence that the environment of early-type galaxies has a significant impact on their X-ray properties (e.g., Brown & Bregman 1998; Helsdon et al. 2001), with isolated ellipticals having the lowest L_X/L_B ratios and the central dominant ellipticals in groups having the highest L_X/L_B ratios. Brighenti & Mathews (1998) have suggested that much of the hot gas in ellipticals arises from the accretion of circumgalactic gas. While the central galaxy in Hydra A is a rather extreme example, since it is located at the center of a rich cluster, it is useful to compare the gaseous properties within the central 10 kpc with that found in samples of more isolated ellipticals. The deprojected X-ray luminosity of the hot gas within the central 10 kpc is $L(0.5-2.0 \text{ keV}) = 8.6 \times 10^{42} \text{ ergs s}^{-1}$ (excluding the central point source). Compared with the samples of ellipticals in Brown & Bregman and Helsdon et al., the central galaxy in Hydra A is 100 times overluminous in X-rays for its observed L_B . The deprojected gas temperature within the central 10 kpc is 2.7 keV which is approximately 3 times higher than a typical elliptical. The higher temperature cannot result from heating by supernovae. Based on the calculations above, the production of the excess Si mass via SNe would only heat the gas by 0.6 keV per particle. Thus, the higher gas temperature must result from gravitational infall and compression. This is consistent with the result noted above that the hot gas in the central galaxy cannot be produced by stellar mass loss and must originate from the intracluster medium.

7. SUMMARY

We have presented a detailed analysis of the *Chandra* observations of the Hydra A cluster. The high spatial resolution of *Chandra* and the good photon statistics in these observations have allowed us to study the distribution of gas temperature, multiphase gas, and metals in the cluster core on scales of 10–20 kpc. These high-resolution observations present significant challenges to the conventional cooling flow scenario and the processes by which the hot gas in clusters is enriched with heavy elements.

The large-scale abundance ratio of heavy elements in the Hydra A cluster determined from the *Chandra* observations indicates that the gas was initially enriched by a predominance of SNe II, in agreement with earlier *ASCA* observations. The uniform distribution of Fe and Si beyond the central 100 kpc suggests that this enrichment occurred prior to the last major merger, during which any preexisting gradients would have been erased. Within the central 100 kpc, the observed excess of Fe and Si has a solar abundance ratio.

The ACIS-S spectra do not show any statistically significant evidence for multiphase gas beyond the central 30 kpc, even through the cooling time of the gas at this radius is

only 1 Gyr. Within 30 kpc, the spectroscopic mass deposition rate is roughly consistent with the observed star formation rate, but is more than an order of magnitude less than the morphological mass accretion rate at the same radius. The gas is also isentropic within the central 30 kpc. Based on these observations, we propose a scenario in which a small amount of the hot gas is able to cool and accrete onto the central black hole in Hydra A and trigger the formation of a radio jet which mechanically heats the gas via strong shocks. We show that the collective effect of numerous weak shocks is energetically insufficient to have a major impact on the cooling flow. The boost in gas entropy after the passage of a strong shock makes the central gas convectively unstable and produces buoyant bubbles of hot gas. Averaged over many outbursts, the mass accretion rate in the outer cooling flow region can be nearly balanced by the mass outflow rate in the expanding bubbles. This scenario makes definitive predictions that can be tested in the near

future with *Chandra* and *XMM*. About 10% of cooling flow clusters should show evidence of strong shocks. By the end of AO-2, *Chandra* will have observed approximately 100 groups and clusters so this prediction can be easily tested. In addition, the greater throughput of *XMM-Newton* may permit the spectroscopic detection of hot buoyant bubbles in the outer parts of cooling flows. Based on our results, it is very likely that some there will be some fundamental changes in the cooling flow scenario over the next few years as data from *Chandra* and *XMM-Newton* continue to accumulate.

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